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Far-infrared constraints on the contamination by dust-obscured galaxies of high- z dropout searches

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ABSTRACT

The spectral energy distributions (SED) of dusty galaxies at intermediate redshift may look similar to very high-redshift galaxies in the optical/near infrared (NIR) domain. This can lead to the contamination of high-redshift galaxy searches based on broad-band optical/NIR photometry by lower redshift dusty galaxies because both kind of galaxies cannot be distinguished. The contamination rate could be as high as 50%. This work shows how the far-infrared (FIR) domain can help to recognize likely low- z interlopers in an optical/NIR search for high- z galaxies. We analyze the FIR SEDs of two galaxies that are proposed to be very high-redshift ($z > 7$) dropout candidates based on deep Hawk-I/VLT observations. The FIR SEDs are sampled with PACS/*Herschel* at 100 and 160 μm , with SPIRE/*Herschel* at 250, 350 and 500 μm and with LABOCA/APEX at 870 μm . We find that redshifts > 7 would imply extreme FIR SEDs (with dust temperatures > 100 K and FIR luminosities $> 10^{13} L_{\odot}$). At $z \sim 2$, instead, the SEDs of both sources would be compatible with those of typical ultra luminous infrared galaxies or submillimeter galaxies. Considering all available data for these sources from visible to FIR we re-estimate the redshifts and find $z \sim 1.6$ – 2.5 . Owing to the strong spectral breaks observed in these galaxies, standard templates from the literature fail to reproduce the visible-to-near-IR part of the SEDs even when additional extinction is included. These sources strongly resemble dust-obscured galaxies selected in *Spitzer* observations with extreme visible-to-FIR colors, and the galaxy GN10 at $z = 4$. Galaxies with similar SEDs could contaminate other high-redshift surveys.

Key words. galaxies: distances and redshifts – dust, extinction – gravitational lensing: weak – galaxies: high-redshift

1. Introduction

Observing galaxies up to very high-redshifts allows us to study directly the formation and evolution of structures in the expanding Universe. Finding galaxies at ever higher redshifts has therefore become one of the main areas of extragalactic astronomy. The most common technique is to use known broad features in the spectral energy distributions (SEDs) of galaxies to identify high-redshift sources in deep optical and near-infrared (NIR) multi-band observations. In particular the Lyman break is widely used to select sources by redshift, noting their disappearance in bands below a given wavelength, the so-called dropout technique

(Steidel et al. 1996). With this technique and state-of-the-art telescopes and instruments it is now possible to select sources that are good candidates for being at the end or within the epoch of reionization (Richard et al. 2006; Zheng et al. 2009; McLure et al. 2010; Wilkins et al. 2010; Oesch et al. 2010; Bouwens et al. 2010b,a).

Low-redshift galaxies, however, can have very steep SEDs that resemble a break in the UV/optical/NIR. This can lead to contamination of the dropout selection of very high- z galaxies, and consequently to erroneous estimates of the star-formation rate density, stellar masses, and others, although these effects

are currently difficult to quantify. Objects such as these have been found and discussed by several authors (see e.g. Dickinson et al. 2000; Mobasher et al. 2005; Schaerer et al. 2007; Dunlop et al. 2007; Chary et al. 2007; Capak et al. 2011). Confirming the photometric redshifts of high- z galaxies by identifying spectral lines is challenging because the sources are generally too faint for spectroscopic follow-up observations or because they may intrinsically lack Ly α emission (but see Vanzella et al. 2011).

The recent developments of space far-infrared (FIR) instrumentation offer new perspectives in this domain. In particular, with the advent of the *Herschel* Space Observatory it is now possible to sample the FIR part of the SEDs, where the thermal dust emission dominates. The shape of the FIR SED universally looks like a broad bump and can be used to further constrain the optical/NIR photometric redshifts. Although the wavelength of the FIR SED peak also depends on the dust temperature, the limited range of average temperatures observed so far in galaxies (between 20 and 60 K averaged over the entire galaxies, see, e.g., Kovács et al. 2006; Magnelli et al. 2010; Magdis et al. 2010; Wardlow et al. 2011) can be used as a prior and makes it possible to distinguish between intermediate ($z < 3$) and very high-redshifts ($z > 6$).

Far-infrared observations of high- z candidates are also essential to characterize their star forming and dust properties and thus interpret correctly their contribution to the cosmic history of star-formation and reionization.

Recently Laporte et al. (2011) identified ten $z > 7$ candidates in the field of the cluster Abell 2667 using photometric dropout criteria based on deep observations with HAWK-I on the ESO Very Large Telescope (VLT). Comparing their results to other studies and in particular to the WIRCAM Ultra Deep Field Survey (WUDS; Pello et al. in prep.), which is based on deeper optical observations blueward of the I -band, they estimated that 50–75% of these candidates could in fact be lower redshift interlopers. Here, we study two galaxies of this sample that are clearly detected by *Herschel*, namely the sources named “z1” and “Y5”. Our goal is to determine whether they could be interlopers and to understand their nature. The redshift probability distributions of these two sources derived by Laporte et al. (2011) from SED fitting to deep optical/NIR photometry show a prominent peak at $z = 7.6$ and 8.6 respectively. However, a secondary peak at lower redshift around $z \sim 2$ indicates that they could be interlopers as well. Laporte et al. (2011) also noted that the $24\mu\text{m}$ detection of z1 with MIPS/*Spitzer* (Y5 is outside the *Spitzer* map) seems difficult to reconcile with the high- z solution. We use new *Herschel* and LABOCA observations of Abell 2667 to reconstruct the FIR part of their SEDs. We can thus further constrain their redshifts and study their physical properties.

The layout of the article is as follows: Sect. 2 gives a presentation of the observations and data analysis. In Sect. 3 the FIR part of the SEDs is analyzed. In Sect. 4 the complete SEDs are used to estimate the redshifts and discuss the physical properties of the two galaxies. In Sect. 5 we compare the two galaxies to other similar galaxies found in the literature. Section 6 gives the conclusions. We assume a Λ -cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$.

2. Observations and data analysis

2.1. Observations and reduction

Herschel observations were obtained in the framework of the *Herschel* Lensing Survey (HLS) described by Egami et al. (2010). They include PACS data at 100 and $160\mu\text{m}$,

and SPIRE observations at 250, 350, and $500\mu\text{m}$. The data reduction was made with the HIPE software as described by Rex et al. (2010) and Rawle et al. (2010).

The large APEX Bolometer Camera (LABOCA Siringo et al. 2009) is a bolometer array operating at $870\mu\text{m}$ that is mounted on the APEX telescope in the desert of Atacama, Chile (Güsten et al. 2006). The LABOCA observations were conducted during the summer 2010. The cluster was mapped in spiral mode during 30 h, covering a circular field of $\sim 6'$ in radius. The data were reduced with the BoA¹ software. The noise is not uniform over the map and the RMS is in the range 1.1–3.0 mJy, the highest values are reached at the edges of the map.

We also obtained a VLA 1.4 GHz continuum map of Abell 2667 (PI: Ivison) with an RMS of $46\mu\text{Jy}$.

2.2. Analysis

The astrometry of all maps was corrected to align them with the VLT K_s image. All optical dropout sources of Laporte et al. (2011) were inspected in the *Herschel* and LABOCA images. Two of them, z1 and Y5, are detected in several FIR bands. Because IRAC/*Spitzer* and MIPS/*Spitzer* data are available for z1, the source can be followed from one band to the next one by increasing wavelength despite the decreasing resolution. Its identification is therefore robust.

For Y5 there is a larger gap in the SED owing to the lack of data between $8\mu\text{m}$ and $100\mu\text{m}$, and due to the fact that it lies at the noisy edges of the 100 and $160\mu\text{m}$ maps, where it is not detected. However, Y5 is the only source detected at $4.5\mu\text{m}$ within a radius of $3''$ (i.e., $\sim 1/3$ rd of the $250\mu\text{m}$ beam radius) around the $250\mu\text{m}$ peak, its identification with the SPIRE detection is therefore very likely.

The fluxes are measured at the positions of the two galaxies by PSF fitting in apertures with a radius equal to $FWHM/3$, where $FWHM$ is the PSF full width at half maximum, i.e., $5.6''$, $11.3''$, $18.1''$, $24.9''$, $36.6''$ and $22.5''$ from 100 to $870\mu\text{m}$. The last (LABOCA) $FWHM$ corresponds to the APEX beam convolved by a Gaussian of $12''$. The sources were deblended from the neighboring sources by subtracting PSFs at the positions of the neighbors derived from the $250\mu\text{m}$ map. Observations at these wavelengths with these resolutions are affected by source confusion. As a consequence a measured flux cannot be directly interpreted as the true flux of a single underlying source. A correct treatment of the effect of source confusion on flux measurements (a.k.a. flux “deboosting”) requires a prior knowledge of the source counts toward low fluxes at the given wavelength. We followed the method presented by Crawford et al. (2010) based on a Bayesian analysis. For the prior source counts we extrapolated toward low fluxes the results of Berta et al. (2010) for PACS bands, Oliver et al. (2010) for SPIRE bands and Coppin et al. (2006) for the LABOCA band.

Blending affects z1 photometry at $\lambda \geq 250\mu\text{m}$ and Y5 photometry at $\lambda \geq 500\mu\text{m}$. And the effect of deboosting is small ($< 20\%$), except for the $870\mu\text{m}$ measurement of Y5, which corresponds to a 2.6σ signal and which we chose to consider as a tentative detection. The deboosted flux of Y5 at $870\mu\text{m}$ is $1.8 \pm 1 \text{ mJy}$ for a measured flux of $2.5 \pm 0.95 \text{ mJy}$.

None of the two sources are detected in the VLA map. The measured FIR fluxes of the sources as well as their optical-to-near-IR photometry from Laporte et al. (2011) are listed in Table 1. Thumbnails of the *Herschel* and LABOCA bands centered at the source positions as well as FIR SED fits are

¹ <http://www.apex-telescope.org/bolometer/laboca/boa/>

Table 1. Multi-wavelength SED of $z1$ (Col. 3) and $Y5$ (Col. 4).

Band/instrument	λ_{eff} [μm]	$z1$	$Y5$
I	0.79	$<3.6\text{e-}5$	$<3.6\text{e-}5$
z	0.92	$<1.4\text{e-}4$	$<1.4\text{e-}4$
Y	1.02	$(2.0 \pm 0.3)\text{e-}4$	$<6.3\text{e-}5$
J	1.26	$(1.75 \pm 0.05)\text{e-}3$	$(6.9 \pm 0.5)\text{e-}4$
H	1.63	$(1.33 \pm 0.07)\text{e-}3$	$(9.3 \pm 0.03)\text{e-}4$
K_s	2.15	$(2.29 \pm 0.06)\text{e-}3$	$(1.77 \pm 0.07)\text{e-}3$
IRAC	3.6	$(6.98 \pm 0.06)\text{e-}3$	$(2.91 \pm 0.11)\text{e-}3$
	4.5	$(10.20 \pm 0.09)\text{e-}3$	$(3.73 \pm 0.10)\text{e-}3$
	8.0	$(9.91 \pm 0.96)\text{e-}3$	
MIPS	24	0.340 ± 0.040	
PACS	100	<3.3	<18
	160	6.3 ± 2.0	<30
SPIRE	250	19.4 ± 1.6	45.5 ± 1.6
	350	15.7 ± 1.4	30.3 ± 1.2
	500	7.5 ± 1.7	19.4 ± 2.1
LABOCA	870	<2.5	1.8 ± 1.0
VLA	2.1×10^5	<0.14	<0.14

Notes. Optical and IR photometry (rows 1–10) is taken from Laporte et al. (2011). Upper limits are 3σ . Rows 12–17 give the *Herschel* and LABOCA source flux. When the measured flux (not the estimated flux) is $<2.5\sigma$ the 3σ value is given as an upper limit. No entry indicates the lack of data. All fluxes are given in milli-Jansky.

shown in the Fig. 1. $Y5$ is close to the border of the PACS maps where the noise is higher, hence the high upper limits.

3. Analysis of the FIR SEDs

The following models or templates were fitted to the FIR measurements (cf. Fig. 1):

- a modified black-body SED parameterized as described by Blain et al. (2003), with emissivity fixed to $\beta = 1.5$ and the Wien correction parameter $\alpha = 2.9$. These values are adapted to submillimeter galaxies (SMGs; Chapman et al. 2005) and local ultraluminous infrared galaxies (ULIRGs; Dunne et al. 2000; Blain et al. 2003). The free parameters are the total FIR luminosity, L_{FIR} , defined as the luminosity emitted in the range 8–1000 μm , and the dust temperature, T_d ;
- the 105 galaxy templates built by Chary & Elbaz (2001). The templates are fitted without rescaling²;
- ULIRG templates built by Vega et al. (2008), with a scaling parameter, L_{FIR} ;
- the starburst, Seyfert, and active galactic nuclei (AGN) templates of Polletta et al. (2007), with a scaling parameter, L_{FIR} ;
- templates built by Michałowski et al. (2010a,b) to fit high-redshift galaxies with detected but poorly sampled submm emission, with a scaling parameter, L_{FIR} ;
- the SED fit to the observations of SMM J2135-0102 (Swinbank et al. 2010; Ivison et al. 2010) with a scaling parameter, L_{FIR} .

The fit was performed by finding the maximum likelihood assuming Gaussian probability distributions for the measurements. When there is no detection, the $3\text{-}\sigma$ value is used as a hard upper limit, i.e., the probability is assumed to be uniform in the $[0, 3\sigma]$ interval and zero outside. The redshifts are fixed to the solutions

² We found that this additional scaling parameter was not required to obtain good fits.

derived by Laporte et al. (2011) from the optical/NIR photometry, i.e., $z = 1.8$ and 7.6 for $z1$ and $z = 1.7$ and 8.6 for $Y5$. The corresponding magnification factors are $\mu = 1.12$ and 1.17 for $z1$ and $\mu = 1.04$ and 1.15 for $Y5$. The MIPS/*Spitzer* 24 μm flux of $z1$ ($Y5$ has no 24 μm data available) was taken into account to fit various galaxy templates, but ignored to fit the modified black-body because it is most likely dominated by polycyclic aromatic hydrocarbons (PAHs).

For both sources we find reasonable fits at low-redshift for the modified black-body and the various galaxy templates. A ULIRG template from the Vega et al. (2008) library, a submillimeter-detected galaxy template from the Michałowski et al. (2010a,b) library and a Seyfert template from the (Polletta et al. 2007) library are able to reproduce the 24 μm emission of $z1$. The modified black-body model gives dust temperatures of 34 and 36 K, for $z1$ and $Y5$ respectively, which are typical values for integrated dust temperatures in LIRGs. The infrared luminosities, L_{FIR} , are in the range $(1.2\text{--}1.7) \times 10^{12} L_{\odot}$ and $(3.1\text{--}4.7) \times 10^{12} L_{\odot}$ for $z1$ and $Y5$ respectively. There is a noticeable agreement in L_{FIR} between the modified black-body model and the various templates. These galaxies would therefore be typical ULIRGs/SMGs at $z \sim 2$. This is consistent with the general picture of galaxy evolution now widely observed, i.e., that the contribution of ULIRGs to the cosmic SFR is expected to peak at $z \sim 2$ where it should be comparable to that of the more “normal” galaxies (see e.g. Murphy et al. 2011).

For the high-redshift solutions ($z > 7.5$), instead, the modified black-body requires for both sources very high dust temperatures, i.e., 105 and 129 K for $z1$ and $Y5$, respectively. Such high temperatures averaged over an entire galaxy are extreme³. This can be seen from the impossibility to find any good fit in the different template libraries, which were built from observed galaxies. However, the dust properties of galaxies at such high-redshifts are unknown and dust temperatures above 100 K cannot be ruled out. This would imply that the FIR luminosities are on the order of 0.5 and $1.5 \times 10^{14} L_{\odot}$ for $z1$ and $Y5$, respectively, i.e., both sources would be classified as hyper luminous infrared galaxies (HyLIRG). While the nature of HyLIRGs is still a matter of debate (e.g. Ruiz et al. 2010) and their density at very high-redshift is not well known, they are extreme sources with a lower number density than ULIRGs.

The radio continuum upper limits are too high to constrain the SED fitting. We find that for $Y5$ only at the low-redshift solution ($z \sim 2$) the Chary & Elbaz template is close to the 3σ upper limit.

Thus, in summary, by comparing the FIR photometry to known galaxy SEDs and by taking into account the expected temperature and luminosity range of high-redshift galaxies, the very high-redshift solutions derived from the optical/NIR photometry seem to be less likely than the low-redshift solution. The two sources are most likely typical ULIRGs at $z \sim 2$. This result puts strong constraints on the optical/NIR analysis, which gave a much higher probability to the very high-redshift solution when no prior luminosity function was taken into account. The FIR data alone, however, cannot be used to derive any accurate photometric redshift because of the redshift-temperature degeneracy.

³ To our knowledge there was only one such extreme case reported so far: the host of the lensed quasar APM08279 at $z = 3.9$ (Weiß et al. 2007; Riechers et al. 2009). It requires two dust components at 75 and 220 K and the highest temperatures most likely result from heating by the powerful quasar.

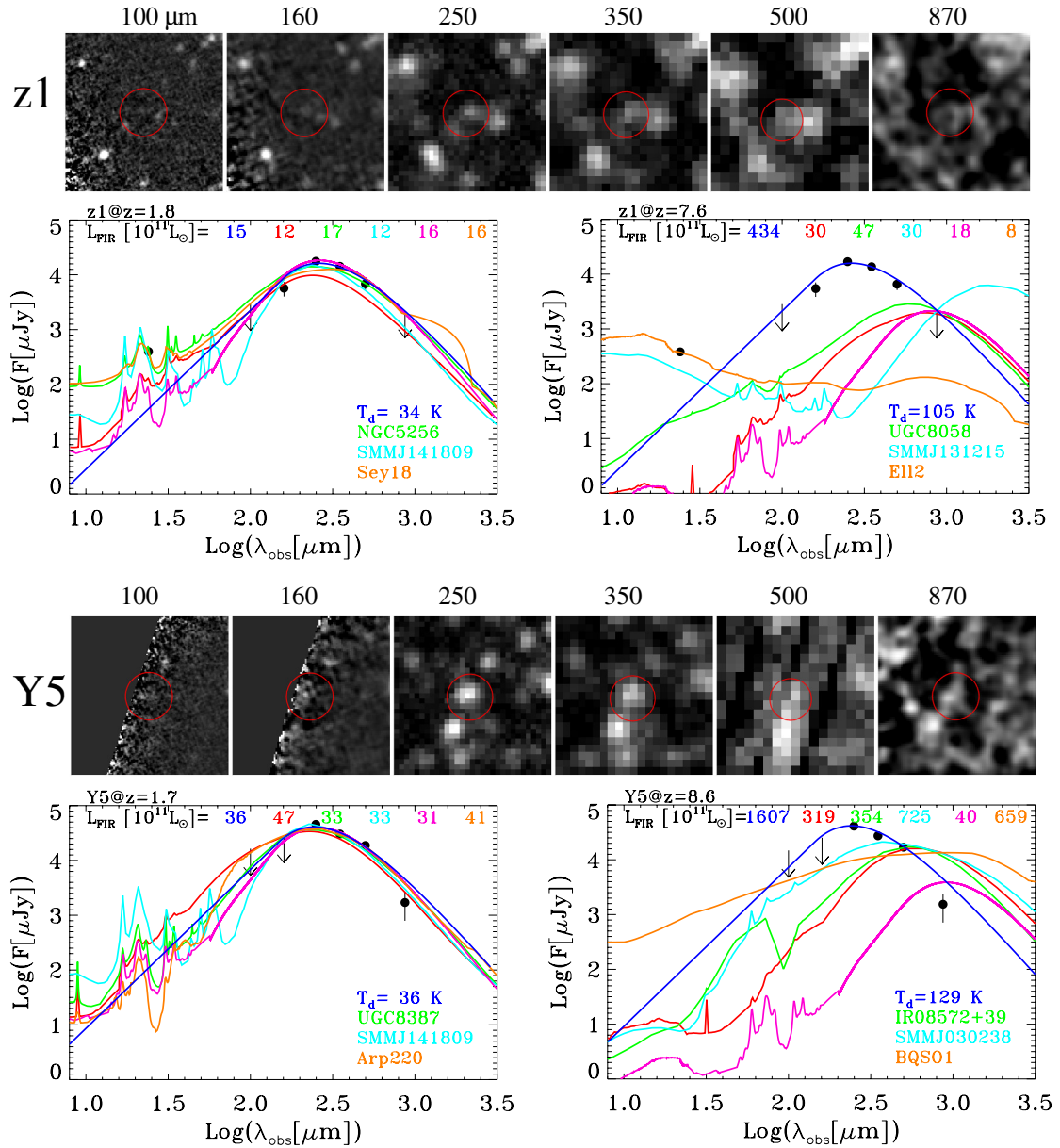


Fig. 1. Top figures show the thumbnails in the five *Herschel* bands (100, 160, 250, 350, 500 μm from left to right) and the LABOCA band (870 μm , the right-most stamp) centered on z1 and the best-fit SEDs for the low-redshift (left) and high-redshift (right) solutions. The bottom figures show the same for the source Y5. The red circle in the thumbnails is meant to guide the eye, it has a radius of $27''$ and is centered on the source position. The blue curves correspond to the best-fit modified black-body SEDs with the parameters written in blue, the magenta curves correspond to the SMMJ2135-0102 model (Swinbank et al. 2010; Ivison et al. 2010), the other curves correspond to the best-fit templates of the Chary & Elbaz (2001) library (red), the Vega et al. (2008) library (green), the Michałowski et al. (2010a,b) library (cyan) and the Polletta et al. (2007) library (orange). The names of the best-fit templates as well as L_{FIR} are written in the figure with the same color codes.

4. Analysis of the complete SED from visible to FIR

We will now examine all data from the visible to the FIR ranges to improve the redshift estimate of our galaxies and to examine the nature and physical properties of these sources.

4.1. Method

To model the SED of the two sources we used a modified version of the *Hyperz* photometric redshift code of Bolzonella et al. (2000) described in Schaerer & de Barros (2009). Non-detections are treated as the usual case 1 of *Hyperz*, i.e., the flux in these filters is set to zero, with an error bar corresponding

to the flux at 1σ level. The basic spectral templates are taken from the Bruzual & Charlot models (Bruzual & Charlot 2003), computed for a variety of star-formation histories and metallicities. Although applicable only to a limited part of the spectrum, we use these templates here to constrain redshift, extinction, and stellar mass in particular. For the Bruzual & Charlot templates we consider variable extinction with A_V up to 8 mag for the Calzetti et al. (2000) attenuation law. We also explored other extinction laws.

The code, initially designed to fit rest-frame UV to near-IR (stellar) emission, can also easily be used to include the thermal mid-IR and beyond. To cover the entire spectral range from the visible to the millimeter domain, and to compare our sources

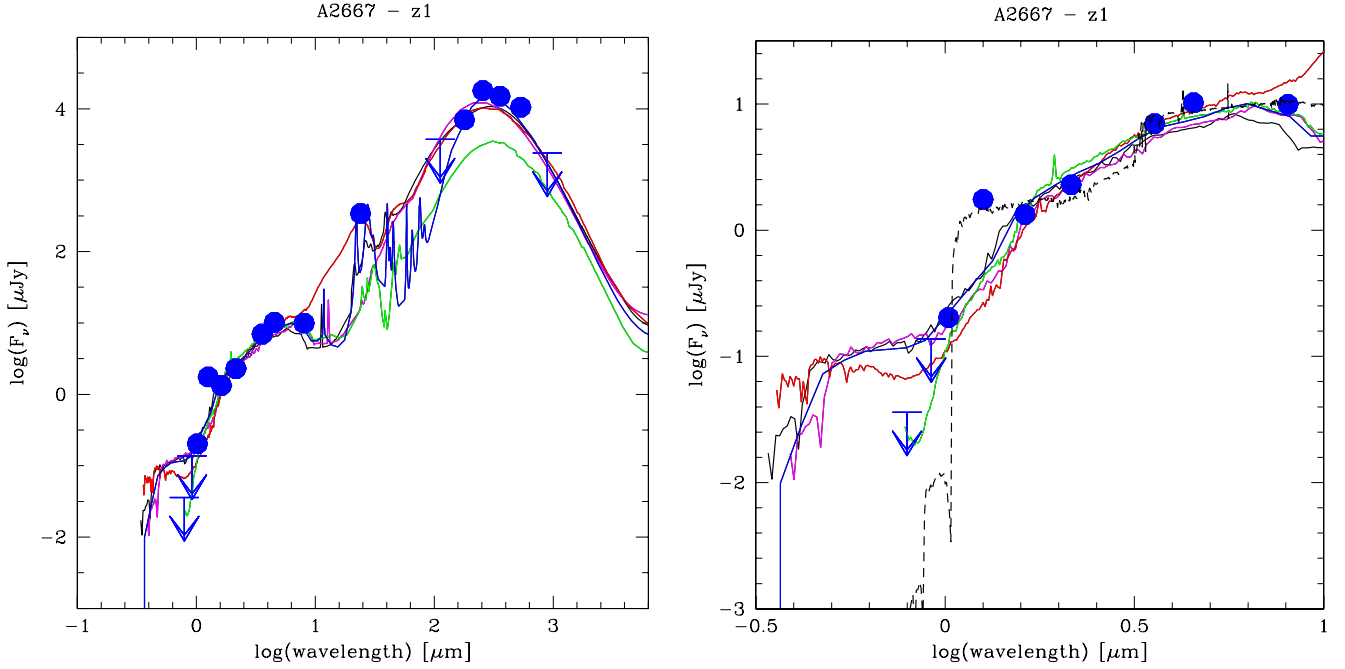


Fig. 2. Fits to the observed SED of source $z1$ (photometry shown by blue symbols, including 3σ upper limits) using different spectral templates: Chary & Elbaz (2001) (black = best-fit template, and magenta = template with maximum IR luminosity), Polletta et al. (2007) (red = global best-fit template, and green = best-fit to visible-near-IR SED excluding the thermal IR), and Michałowski et al. (2010a) (blue). The best-fit SED with the templates of Vega et al. (2008), very similar to the one using Polletta’s templates, is not shown here for simplicity. *Left:* global visible to sub-mm SED. *Right:* zoom on visible to near-IR part of the SED including for comparison also the best-fit SED at high-redshift ($z = 7.5$) from Laporte et al. (2011, dashed line), which is most likely excluded because of our Herschel detections.

with SEDs of very different galaxy types, we compiled a great variety of spectral templates from the GRASIL models of Silva et al. (1998), the library of Chary & Elbaz (2001), Rieke et al. (2009), the starburst, Seyfert, and AGN templates of Polletta et al. (2007), the ULIRG templates of Vega et al. (2008), the sub-mm galaxy templates of Michałowski et al. (2010a), and the model fit to SMMJ2135-0102 (Swinbank et al. 2010; Ivison et al. 2010). Extinction can also be added to these spectral templates; SED fits with and without additional extinction will be discussed below.

We carried out both fits of the entire SED (optical, near-IR, and IR) and fits up to $8\ \mu\text{m}$ only (for the Bruzual & Charlot templates). For each template set the free parameters are redshift and (additional) A_V . Physical parameters such as the infrared luminosity, L_{IR} , defined as the luminosity emitted in the range $2\text{--}1000\ \mu\text{m}$; the IR star-formation rate, SFR; and the stellar mass, are subsequently derived from the best-fit templates. In contrast to the IR fits discussed in Sect. 3 we have no handle on the dust temperature, because this is not a parameter describing the SEDs used here. We also checked that the two independent fitting methods used here and in Sect. 3 give consistent results.

4.2. Photometric redshifts

As discussed in depth by Laporte et al. (2011), the best-fit photometric redshifts of our sources derived from the optical-to-near-IR photometry (up to $8\ \mu\text{m}$) and using standard spectral templates is consistently found at $z > 7$ with a lower probability at low z . This result remains unchanged with the exploration of a wider range of extinction, different attenuation/extinction laws, and template sets used here compared to Laporte et al. (2011). On the other hand, analysis of the IR SED and other arguments clearly favor low redshifts ($z \sim 1.5\text{--}2.5$), as discussed above.

Below we therefore limit ourselves to $z < 4$ and attempt to refine the photometric redshift of the two sources.

4.3. Results for $z1$

Overall the global SED fits for this source are fairly satisfactory, as shown in Fig. 2, albeit with significant discrepancies in the optical domain (cf. below). Several templates (from Polletta et al. 2007; Vega et al. 2008; Michałowski et al. 2010a) also reproduce the $24\ \mu\text{m}$ flux, and the observed 100 and $870\ \mu\text{m}$ fluxes are within $2\text{--}3\sigma$ of the model. Interestingly, the best-fits for both, the Polletta and Michałowski libraries are found with templates for active galaxies.

The best-fit redshift found with these templates is between $z \approx 2.24$ and 2.57 . The resulting IR luminosity is $L_{\text{IR}} \sim (2.6\text{--}3.2) \times 10^{12} L_{\odot}$, the corresponding $SFR \approx 450\text{--}550 M_{\odot} \text{ yr}^{-1}$ using the standard Kennicutt (1998) calibration. Fits to the IR part with the SMMJ2135-0102 template yield $z_{\text{phot}} \approx 2.0$. A somewhat lower redshift of $z \approx 1.7$ is found with the Bruzual & Charlot (2003) templates using the SED up to the IRAC bands. The estimated extinction is $A_V \sim 2.6$, the stellar mass $M_{\star} \sim 6 \times 10^{10} M_{\odot}$ for the same Salpeter IMF as adopted by Kennicutt (1998). However, these values should be taken with caution because the fits are poor. For comparison, one obtains $M_{\star} \sim 3 \times 10^{10} M_{\odot}$ from the absolute H -band magnitude ($M_B \approx -23.0$) using the mass-to-light ratio adopted by Wardlow et al. (2011) for SMGs.

At a more detailed level (see right panel), all spectral templates have some difficulty to reproduce the steep, observed SED between the visible (I , z bands) and the near-IR (Y and J here), and they predict a flux excess in the optical domain. Below we will show that this also holds when variable extinction is added to the empirical templates. The same is also true for all other templates we examined, including the theoretical galaxy

templates of [Bruzual & Charlot \(2003\)](#). This sharp drop is of course the reason why this source was selected as an optical dropout (Y -drop).

4.4. Results for $Y5$

For this source the global fits are less good than for $z1$. This is because $Y5$ shows a higher flux ratio between the thermal-IR and the near-IR than $z1$, whose SED already required templates with extreme IR/near-IR fluxes. For example, the Arp 220 template from [Polletta et al. \(2007\)](#), shown in red, underpredicts the IR flux by a factor $\gtrsim 5$. The only template coming close to the observed IR emission is from the SMG library of [Michałowski et al. \(2010a\)](#) (SMMJ221725.97+001238). With a best-fit redshift of $z_{\text{phot}} \approx 2.15$ this translates to $L_{\text{IR}} = 2.2 \times 10^{12} L_{\odot}$, corresponding to $SFR \approx 380 M_{\odot} \text{ yr}^{-1}$. Fits to the IR part with the SMMJ2135-0102 template yield $z_{\text{phot}} \approx 1.8$. A best-fit redshift of $z \approx 1.95$ is found with [Bruzual & Charlot \(2003\)](#) templates using the SED up to the IRAC bands. The estimated extinction is $A_V \sim 1.6$, the stellar mass $M_{\star} \sim 3 \times 10^{10} M_{\odot}$. However, these values should be taken with caution because the fits in the domain close to the optical are poor. Again, using the absolute H -band magnitude ($M_B \approx -22.3$), one obtains $M_{\star} \sim 2 \times 10^{10} M_{\odot}$ with the assumptions already mentioned above.

Similarly to $z1$, the visible-near-IR drop of the SED (see right panel) is poorly fitted by the spectral templates, predicting that the source should be detectable in the visible (I , z , Y bands in particular), in contrast to our observations. The template that fits this part of the spectrum best is an S0 template from [Polletta et al. \(2007\)](#), shown in green. However, this template underpredicts the IR emission by several orders of magnitudes.

4.5. Possible explanations for the strong SED break/very red spectrum

As already seen, the common, observed spectral templates fail to reproduce the steep, observed SED between the visible (I , z bands) and the near-IR (Y and J here), and they predict a flux excess in the optical domain. What causes the sharp observed decrease of the flux between the near-IR and the optical for these sources? The main difficulty arises because the largest spectral break known in galaxy spectra is the Lyman break, whereas the typically observed Balmer (or “4000 Å”) break is smaller than that of our two galaxies. We have examined various possibilities, but with no convincing answer.

For example, as shown in Fig. 4 for $z1$, adding variable extinction to the empirical templates allows one to diminish their optical flux excess somewhat. For the Chary & Elbaz templates, the best-fit (to the domain shown here) is then with an additional extinction $\Delta A_V = 1.4$ for the Calzetti law. However, the templates are too smooth to reproduce the apparent break. The flux at optical wavelengths can be reduced even more by assuming a steeper attenuation law than the Calzetti law adopted by default. With the SMC extinction law by [Prevot et al. \(1984\)](#) the best-fit is then for $\Delta A_V = 1.2$. Overall, the same difficulty remains, however. Figure 4 also shows the best-fit SED for the [Bruzual & Charlot \(2003\)](#) templates with the SMC law. Although it is the steepest SED between the optical and the H -band, it falls short in flux in the J -band. For completeness we also examined templates from the synthesis models of [Maraston et al. \(2006\)](#). As expected, these models do not yield significantly different fits in the blue part of the rest-frame optical spectrum.

As already mentioned by [Laporte et al. \(2011\)](#), we also attempted to fit the SEDs with our models including nebular lines (see [Schaerer & de Barros 2009](#)). Indeed, in this case the best-fit is found at $z \sim 1.5$ such that the [OIII] $\lambda\lambda$ 4959, 5007 lines, and $H\beta$ boost somewhat the J -band flux, and $H\alpha$ the H -band to a lesser extent, contributing thus to the flux decrement between J and Y . However, this solution also requires a very high attenuation ($A_V \sim 4.0$ for the SMC law) to reproduce the steeply rising SED toward longer wavelengths⁴. Although to the best of our knowledge objects with these red SEDs and strong emission lines are not known, this extreme explanation should be easy to test with spectroscopic observations.

Finally, could composite populations not taken into account by our models help to explain the observed SED? Certainly the theoretical SED models may suffer from this simplification. However, we do not see how this could help to resolve the problem with the large observed spectral break, because a superposition of individual simple stellar populations (not capable of reproducing this observation) can only average out spectral features. We conclude that we have no convincing explanation for the observed sharp drop of the SED of our two sources.

5. Discussion

5.1. Comparison with other objects in the literature

How do our sources compare with other known galaxies and what is their nature?

By design our sources are near-IR-selected, optical dropout sources, i.e., sources with a very red color between the J - and z -band and/or between Y and J . Our sources can therefore be compared to those selected by [Capak et al. \(2011\)](#) from the COSMOS survey. From their Fig. 13 we note that with $(I - J) > 5.4$ and 4.4 and $(z - J) > 3.9$ and 3.1 for $z1$ and $Y5$, both sources show extreme (very red) optical-to-near-IR colors, when compared to other low-redshift galaxies with red $(J - z)$ colors. $z1$ and $Y5$ are also similar to the z -dropout galaxy HUDF-J2 identified by [Mobasher et al. \(2005\)](#) as a $z \approx 6.5$ post-starburst galaxy candidate, but later shown to be most likely at $z \sim 1.8$ – 2.5 interloper by [Schaerer et al. \(2007\)](#); [Dunlop et al. \(2007\)](#); [Chary et al. \(2007\)](#). Although similar in several respects, HUDF-J2 shows a more monotonously rising SED between z , J , and H than our objects exhibiting a “sharper” break. The colors and fluxes of our sources are also very similar to those of the lensed optical dropout galaxy #2 found behind the cluster Abell 1835, identified with the $z = 2.93$ sub-mm galaxy SMMJ14009+0252 (see [Schaerer et al. 2007](#); [Weiß et al. 2009](#), and references therein).

Our sources are obviously also characterized by a high IR-to-optical flux ratio, a criterion, which has been used by various authors. For example, [Rodighiero et al. \(2007\)](#) have studied IRAC 3.6 μm -selected sources undetected in deep optical *HST* images. The $(K-3.6)$ and $(z-3.6)$ colors of $z1$ and $Y5$ are comparable to their sources; the main difference seems to be in $(H - K)$, where our sources are bluer than those of [Rodighiero et al. \(2007\)](#). $z1$ and $Y5$ appear to be related to the very dusty $z \sim 2$ – 3 sources from this study. [Fiore et al. \(2008\)](#) and [Dey et al. \(2008\)](#) have examined 24 μm selected sources with very red colors between 24 μm and the R band ($S(24)/S(R) \gtrsim 1000$). They concluded that the bulk of these sources are very luminous strongly dust-obscured galaxies (referred to as DOGs) at $z \sim 2$, powered by AGN and/or by starbursts. Using the I -band

⁴ Our model assumes identical attenuation for the continuum and nebular lines, as in [Schaerer & de Barros \(2009\)](#).

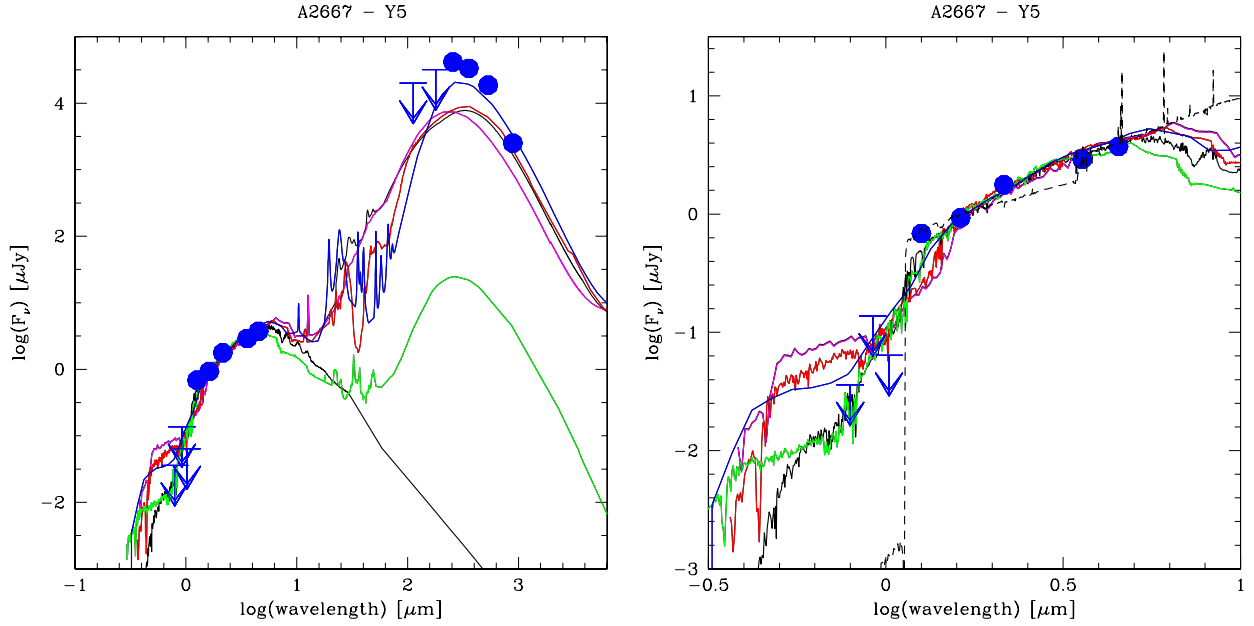


Fig. 3. Same as Fig. 2 for Y5. The best-fit SED at high-redshift (black dashed line) is for $z = 8.28$.

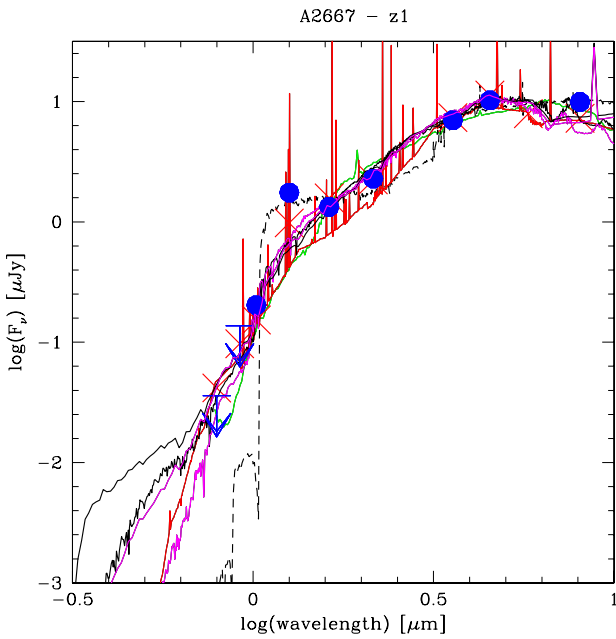


Fig. 4. Observed visible to near-IR part of the SED of $z1$ (blue circles) and best-fit models to this spectral range: using the Polletta et al. (2007) templates and no additional extinction (green line), black lines: using the Chary & Elbaz (2001) and Bruzual & Charlot (2003) templates with variable attenuation described by the Calzetti law, magenta: same as black lines, but adoption the SMC extinction law, red line (continuum) and red crosses (synthetic flux in filters): Bruzual & Charlot (2003) templates with nebular emission and SMC law, black dashed: high-redshift ($z = 7.5$) template from Laporte et al. (2011, dashed line). In all cases the best-fit redshift is $z \approx 1.5$ – 1.7 , except for Polletta’s templates.

as a proxy for R , we obtain a flux ratio $S(24)/S(R) \gtrsim 20\,000$ for $z1$, an extremely high flux ratio compared to the other samples. From the SED of Y5 (cf. Fig. 3), we also expect this galaxy to show a high $24\,\mu\text{m}$ to optical ratio. According to the source density from Dey et al. (2008), we would have expected \sim four strongly dust-obscured galaxies down to $\sim 0.3\,\text{mJy}$ at $24\,\mu\text{m}$ in our $45\,\text{arcmin}^2$ field. Because the depth of our MIPS observations is similar, this value is comparable to our source density,

although our selection is different. Pope et al. (2008) have also compared DOGs and sub-mm galaxies (SMGs) in the same $R - K - 24\,\mu\text{m}$ colors, showing that $\sim 30\%$ of SMGs satisfy the DOG criteria, the remainder showing less extreme (i.e. bluer) colors. This confirms that the SEDs of our sources are comparable to a subset of SMGs with the most extreme optical to IR/sub-mm colors, as already seen above (Sect. 4). Our optical data are not deep enough to ascertain whether $z1$ and Y5 fulfill the usual criteria for extremely red objects (EROs), $(R - K) \gtrsim 5.6$ in Vega magnitudes, at least as estimated from $(I - K_s)$.

Among known sub-mm galaxies, one source, GN10 or GOODS 850-5, stands out as having particularly extreme IR/sub-mm to visible/near-IR properties, similar to our two galaxies. Indeed, this source is undetected down to $\sim 0.01\,\mu\text{Jy}$ (1σ) in the visible, shows fluxes of ~ 1 – $5\,\mu\text{Jy}$ in the IRAC bands (3.6 – $8\,\mu\text{m}$), and peaks at ~ 10 – $20\,\text{mJy}$ around $1\,\text{mm}$ (Wang et al. 2004, 2009; Daddi et al. 2009), quite comparable to $z1$ and Y5. However, GN10 remains undetected even at JHK (Wang et al. 2009), which can be explained by its higher redshift ($z \sim 4$), recently confirmed from CO spectroscopy (Daddi et al. 2009). The observed spectral break of GN10 found between 3.6 and $2.2\,\mu\text{m}$ and other considerations (Wang et al. 2009; Daddi et al. 2009) suggests a very high attenuation of $A_V \gtrsim 4.5$ – 5 for this source, or at least for the star-forming part of it, if hosting multiple components. If we assume constant star-formation as Daddi et al. (2009) for their SED modeling, we would infer $A_V \sim 3$ (7.8 mag for $z1$ (Y5). Comparing the infrared-derived SFR with the upper limits in the rest-frame UV domain, we can also estimate the attenuation of our sources. Adopting the I -band flux as a constraint for the UV flux at ~ 2300 – $2500\,\text{\AA}$ and using the Kennicutt (1998) calibration, we obtain $A_V \gtrsim 4$ mag for both sources.

Sub-mm galaxies are also known to exhibit very strong attenuation. For example, sources with Balmer decrement measurements indicate $A_V \sim 1$ – 3 (Swinbank et al. 2004; Takata et al. 2006), and from SED fits (Swinbank et al. 2004) estimate $A_V = 3.0 \pm 1.0$ for their sample. Wardlow et al. (2011) find $A_V = 2.6 \pm 0.2$ from the median SED of sub-mm galaxies, but more extreme attenuations are found within the sub-mm

galaxy samples (see e.g. [Smail et al. 2004](#)). Also, the extremely red object, sub-mm detected galaxy HR10 studied by [Dey et al. \(1999\)](#) shows $A_V \sim 4.5$, as inferred from comparison of the IR and H α star-formation rate. In short, although higher than the typical/median value of sub-mm galaxies, the attenuation of the sources z1 and Y5 is similar to that of some sub-mm galaxies, such as GN10 at $z = 4.04$ and others at lower z . Our sources are also somewhat fainter, both in the rest-frame near-IR and in the IR, than the typical sub-mm galaxies (cf. [Wardlow et al. 2011](#)). Finally, our sources stand out by their large spectral break, which – to the best of our knowledge – is unusual among intermediate-redshift sources.

5.2. The other high- z candidates of the survey

Based on our FIR detections we have identified two potential interlopers among the ten high- z candidates discovered by [Laporte et al. \(2011\)](#). Most of the other candidates are in crowded regions where several sources emit in the FIR and are blended with each other, making any FIR measurement impossible. Two other candidates only seem to be clean from any contamination in the MIPS, PACS and SPIRE maps, namely Y3 and Y4. They remain undetected in all bands. However, the FIR upper limits obtained do not allow us to distinguish between low and high-redshift. On the other hand we can rule out that these sources are as extreme as z1 and Y5 in their IR/sub-mm to near-IR flux ratio, because they should otherwise clearly be detected in our Herschel images.

6. Conclusions

Analyzing the FIR SED of two high-redshift dropout candidates we find that both galaxies are likely at $z \sim 2$ rather than $z > 7$. From the FIR point of view alone, both galaxies could be similar to ULIRGs or SMGs, which are common at $z \sim 2$. At $z > 7$ the SEDs would imply extreme dust temperatures and luminosities. Fitting the global SEDs considering all available data from visible to submm we estimate $z \sim 1.6$ – 2.5 .

However, the optical/NIR part of both objects remains difficult to understand if at $z \sim 2$. They show a very strong and well-defined spectral break (presumably the Balmer break), unusual among intermediate redshift sources. The source z1 is extreme for both IR/visible and near-IR/optical colors. The source Y5 has a somewhat smaller drop between the near-IR/optical domain.

We have examined several possible explanations for the extreme colors of these galaxies, but none of them is entirely satisfactory. More observations are required to understand their nature. Once a spectroscopic confirmation of their redshift is obtained we will be able to create new SED templates.

Other extreme sources (e.g. GN10 [Wang et al. 2004](#); [Lutz et al. 2005](#); [Wang et al. 2009](#); [Daddi et al. 2009](#)) were found from MIPS observations and submm searches. There could therefore be two complementary paths leading to similar, extreme sources.

Although spectroscopic confirmation is still required and not all possible interlopers may be detected in the FIR, this work shows that FIR observations can be very helpful to constrain the contamination of high- z dropout searches by lower redshift galaxies.

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References

- Berta, S., Magnelli, B., Lutz, D., et al. 2010, *A&A*, 518, L30
 Blain, A. W., Barnard, V. E., & Chapman, S. C. 2003, *MNRAS*, 338, 733
 Bolzonella, M., Miralles, J.-M., & Pelló, R. 2000, *A&A*, 363, 476
 Bouwens, R. J., Illingworth, G. D., González, V., et al. 2010a, *ApJ*, 725, 1587
 Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2010b, *ApJ*, 709, L133
 Bruzual, G., & Charlot, S. 2003, *MNRAS*, 344, 1000
 Calzetti, D., Armus, L., Bohlin, R. C., et al. 2000, *ApJ*, 533, 682
 Capak, P., Mobasher, B., Scoville, N. Z., et al. 2011, *ApJ*, 730, 68
 Chapman, S. C., Blain, A. W., Smail, I., & Ivison, R. J. 2005, *ApJ*, 622, 772
 Chary, R., & Elbaz, D. 2001, *ApJ*, 556, 562
 Chary, R.-R., Teplitz, H. I., Dickinson, M. E., et al. 2007, *ApJ*, 665, 257
 Coppin, K., Chapin, E. L., Mortier, A. M. J., et al. 2006, *MNRAS*, 372, 1621
 Crawford, T. M., Switzer, E. R., Holzapfel, W. L., et al. 2010, *ApJ*, 718, 513
 Daddi, E., Dannerbauer, H., Krips, M., et al. 2009, *ApJ*, 695, L176
 Dey, A., Graham, J. R., Ivison, R. J., et al. 1999, *ApJ*, 519, 610
 Dey, A., Soifer, B. T., Desai, V., et al. 2008, *ApJ*, 677, 943
 Dickinson, M., Hanley, C., Elston, R., et al. 2000, *ApJ*, 531, 624
 Dunlop, J. S., Cirasuolo, M., & McLure, R. J. 2007, *MNRAS*, 376, 1054
 Dunne, L., Eales, S., Edmunds, M., et al. 2000, *MNRAS*, 315, 115
 Egami, E., Rex, M., Rawle, T. D., et al. 2010, *A&A*, 518, L12
 Fiore, F., Grazian, A., Santini, P., et al. 2008, *ApJ*, 672, 94
 Güsten, R., Nyman, L. Å., Schilke, P., et al. 2006, *A&A*, 454, L13
 Ivison, R. J., Swinbank, A. M., Swinyard, B., et al. 2010, *A&A*, 518, L35
 Kennicutt, Jr., R. C. 1998, *ARA&A*, 36, 189
 Kovács, A., Chapman, S. C., Dowell, C. D., et al. 2006, *ApJ*, 650, 592
 Laporte, N., Pelló, R., Schaerer, D., et al. 2011, *A&A*, 531, A74
 Lutz, D., Yan, L., Armus, L., et al. 2005, *ApJ*, 632, L13
 Magdis, G. E., Elbaz, D., Hwang, H. S., et al. 2010, *MNRAS*, 409, 22
 Magnelli, B., Lutz, D., Berta, S., et al. 2010, *A&A*, 518, L28
 Maraston, C., Daddi, E., Renzini, A., et al. 2006, *ApJ*, 652, 85
 McLure, R. J., Dunlop, J. S., Cirasuolo, M., et al. 2010, *MNRAS*, 403, 960
 Michałowski, M., Hjorth, J., & Watson, D. 2010a, *A&A*, 514, A67
 Michałowski, M. J., Watson, D., & Hjorth, J. 2010b, *ApJ*, 712, 942
 Mobasher, B., Dickinson, M., Ferguson, H. C., et al. 2005, *ApJ*, 635, 832
 Murphy, E. J., Chary, R., Dickinson, M., et al. 2011, *ApJ*, 732, 126
 Oesch, P. A., Bouwens, R. J., Illingworth, G. D., et al. 2010, *ApJ*, 709, L16
 Oliver, S. J., Wang, L., Smith, A. J., et al. 2010, *A&A*, 518, L21
 Polletta, M., Tاجر, M., Maraschi, L., et al. 2007, *ApJ*, 81
 Pope, A., Bussmann, R. S., Dey, A., et al. 2008, *ApJ*, 689, 127
 Prevot, M. L., Lequeux, J., Prevot, L., Maurice, E., & Rocca-Volmerange, B. 1984, *A&A*, 132, 389
 Rawle, T. D., Chung, S. M., Fadda, D., et al. 2010, *A&A*, 518, L14
 Rex, M., Rawle, T. D., Egami, E., et al. 2010, *A&A*, 518, L13
 Richard, J., Pelló, R., Schaerer, D., Le Borgne, J., & Kneib, J. 2006, *A&A*, 456, 861
 Riechers, D. A., Walter, F., Carilli, C. L., & Lewis, G. F. 2009, *ApJ*, 690, 463
 Rieke, G. H., Alonso-Herrero, A., Weiner, B. J., et al. 2009, *ApJ*, 692, 556
 Rodighiero, G., Cimatti, A., Franceschini, A., et al. 2007, *A&A*, 470, 21
 Ruiz, A., Miniutti, G., Panessa, F., & Carrera, F. J. 2010, *A&A*, 515, A99
 Schaerer, D., & de Barros, S. 2009, *A&A*, 502, 423
 Schaerer, D., Hempel, A., Egami, E., et al. 2007, *A&A*, 476, 97
 Silva, L., Granato, G. L., Bressan, A., & Danese, L. 1998, *ApJ*, 509, 103
 Siringo, G., Kreysa, E., Kovács, A., et al. 2009, *A&A*, 497, 945
 Smail, I., Chapman, S. C., Blain, A. W., & Ivison, R. J. 2004, *ApJ*, 616, 71
 Steidel, C. C., Giavalisco, M., Pettini, M., Dickinson, M., & Adelberger, K. L. 1996, *ApJ*, 462, L17
 Swinbank, A. M., Smail, I., Chapman, S. C., et al. 2004, *ApJ*, 617, 64
 Swinbank, A. M., Smail, I., Longmore, S., et al. 2010, *Nature*, 464, 733
 Takata, T., Sekiguchi, K., Smail, I., et al. 2006, *ApJ*, 651, 713
 Vanzella, E., Pentericci, L., Fontana, A., et al. 2011, *ApJ*, 730, L35
 Vega, O., Clemens, M. S., Bressan, A., et al. 2008, *A&A*, 484, 631
 Wang, W., Cowie, L. L., & Barger, A. J. 2004, *ApJ*, 613, 655
 Wang, W., Barger, A. J., & Cowie, L. L. 2009, *ApJ*, 690, 319
 Wardlow, J. L., Smail, I., Coppin, K. E. K., et al. 2011, *MNRAS*, 917
 Weiß, A., Downes, D., Neri, R., et al. 2007, *A&A*, 467, 955
 Weiß, A., Ivison, R. J., Downes, D., et al. 2009, *ApJ*, 705, L45
 Wilkins, S. M., Bunker, A. J., Ellis, R. S., et al. 2010, *MNRAS*, 403, 938
 Zheng, W., Bradley, L. D., Bouwens, R. J., et al. 2009, *ApJ*, 697, 1907